



## Cluster observations in the magnetosheath: anisotropies and intensity of the turbulence at electron scales

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The intensities  $\delta B^2$  and  $\delta E^2$  of the magnetic and electric fluctuations measured by STAFF-SA in the magnetosheath (between 8 Hz and 4 kHz) depend strongly on the angle  $\Theta_{BV}$  between the magnetic field  $B$  and the flow velocity  $V$ . This is due to the Doppler effect and implies i) that  $\delta B^2$  and  $\delta E^2$ , in the observed range of wave number  $k$ , have power law spectra  $k^{-\nu}$  ii) that the wave vectors are mostly perpendicular to  $B$  at the electromagnetic (e.m.) scales  $kc/\omega_{pe} \simeq 0.3$  to 30, and mostly parallel to  $B$  at the electrostatic (e.s.) scales  $k\lambda_{De} \simeq 0.1$  to 1 ( $c/\omega_{pe}$  is the electron inertial length,  $\lambda_{De}$  is the Debye length). Parameters other than  $\Theta_{BV}$  could play a part in the intensities  $\delta B^2$  and  $\delta E^2$ , observed during about 23 hours on four different days. For instance, at a given frequency in the spacecraft frame,  $\delta B^2$  and  $\delta E^2$  increase when the solar wind dynamic pressure  $P_{DYN}$  increases. But this increase is only due to the Doppler shift, which is larger when  $P_{DYN}$  is larger. In the plasma rest frame, the intensity of the e.m. and e.s. fluctuations do not depend on  $P_{DYN}$ , nor on the bow shock angle, nor on the magnetosheath  $\beta_{p\parallel}$  or the  $T_e/T_p$  temperature ratio. As  $\delta B^2$  and  $\delta E^2$  depend on  $\Theta_{BV}$  in the spacecraft frame, we look at the distribution of  $\Theta_{BV}$  in the magnetosheath: numerical 3D MHD simulations show where are the regions where  $\Theta_{BV}$  reaches  $90^\circ$ . The most intense e.m. waves and the least intense e.s. waves are observed in these regions, at a given frequency.