

Generation and transport of liquid water in the high-pressure ice layers of Ganymede and Titan

Klára Kalousová¹, Christophe Sotin²

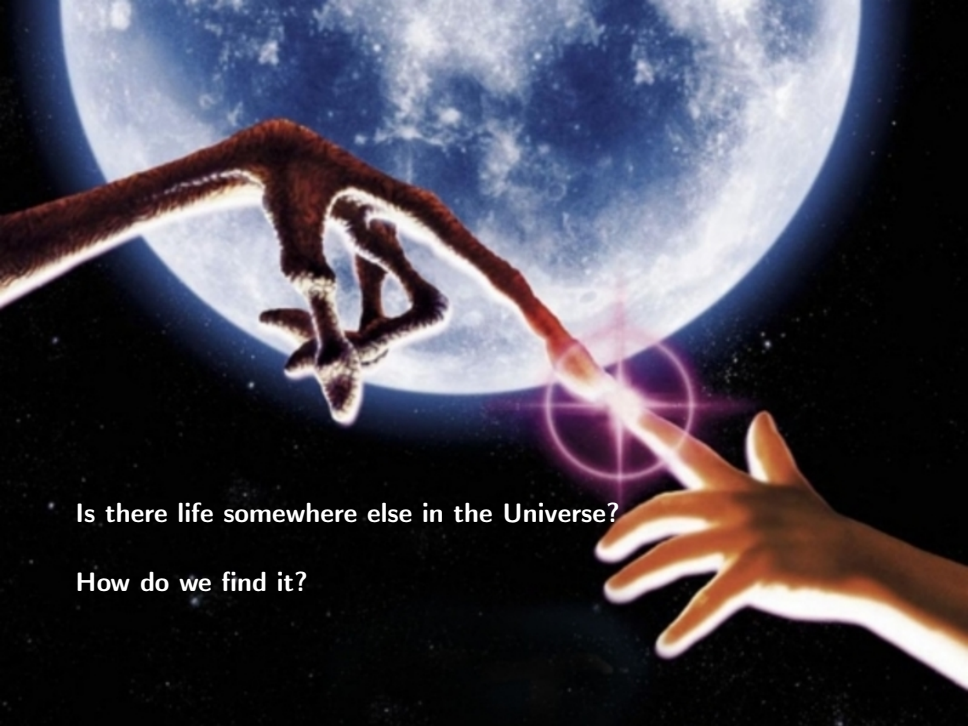
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of Mantle and Lithosphere Dynamics,
La Certosa di Pontignano, Italy,
August 26, 2019

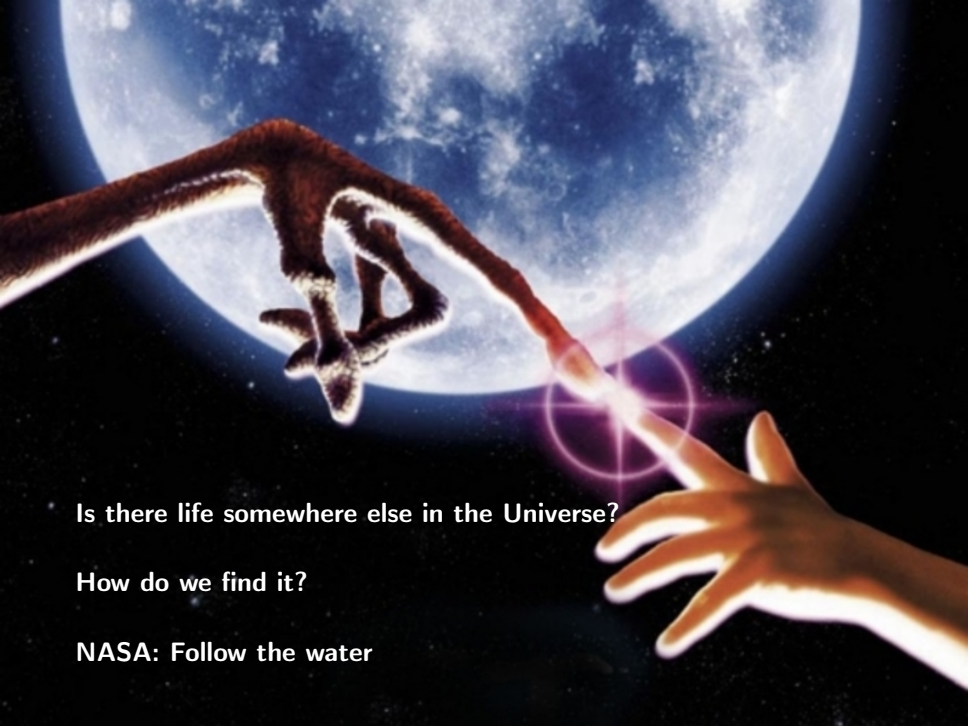


Is there life somewhere else in the Universe?



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How do we find it?



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NASA: Follow the water

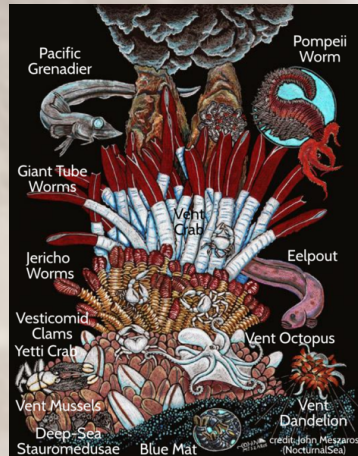
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- ▶ life might have started at the bottom of Earth's oceans



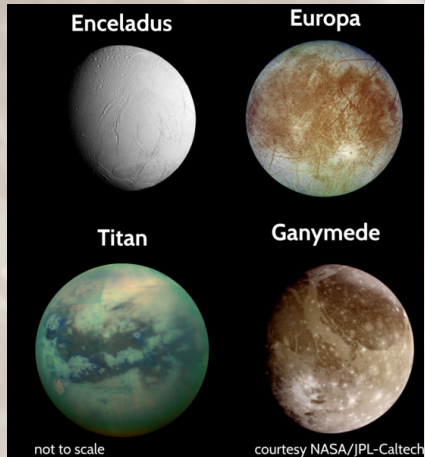
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- ▶ water-rock interactions → rich communities of organisms



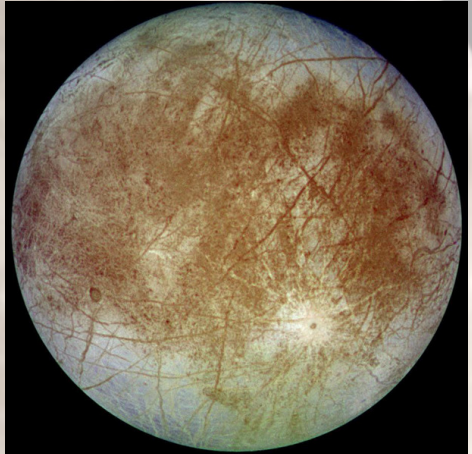
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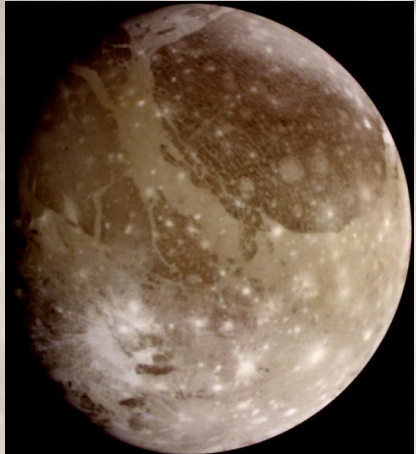
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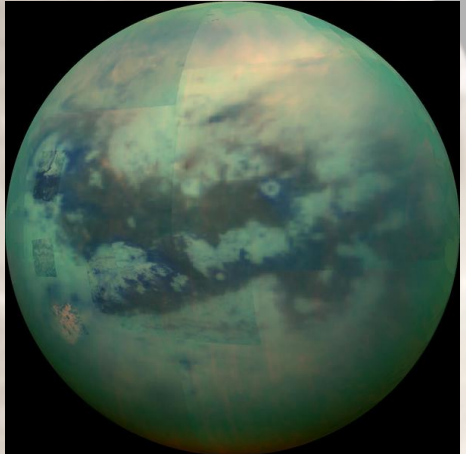
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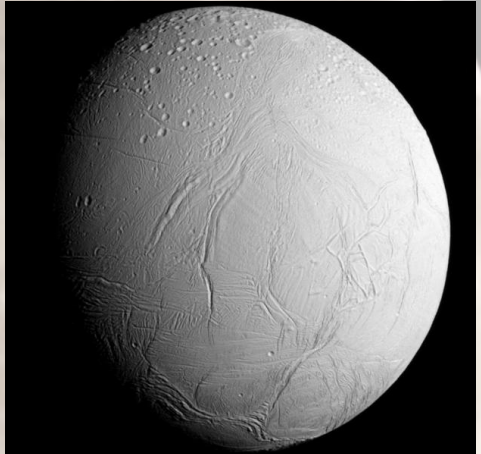
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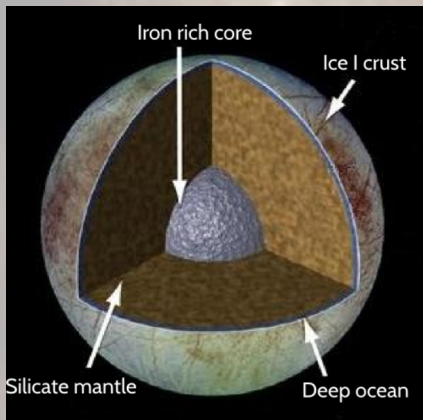
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 - ▶ **Enceladus:**
 - geysers
 - libration of outer layer



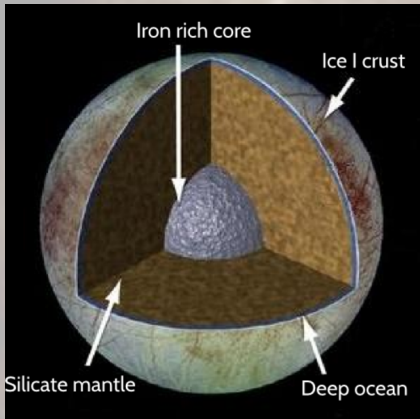
Europa & Enceladus

- ▶ direct contact of silicates with the ocean
- ▶ conditions similar to the Earth oceans - emergence of life?



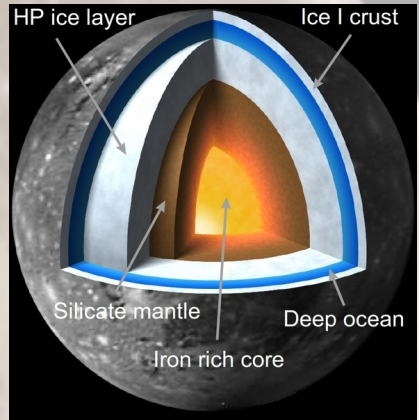
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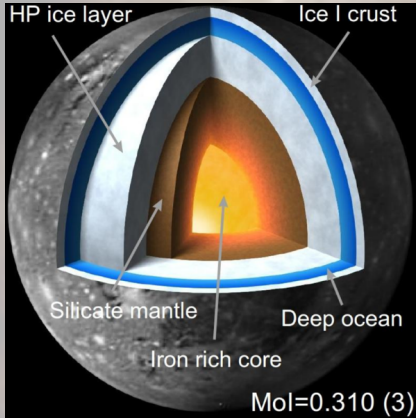
Ganymede & Titan

- ▶ thick HP ice layer between silicates and the ocean
- ▶ is water & material exchange possible?



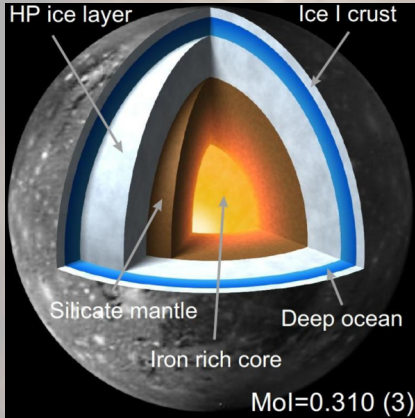
Ganymede

- ▶ full differentiation:
 $Mol=0.310$
- ▶ tenuous O_2 exosphere
- ▶ intrinsic magnetic field



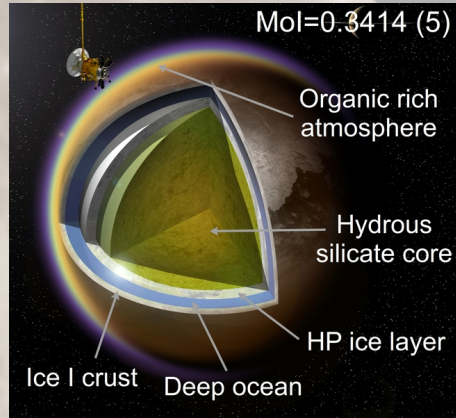
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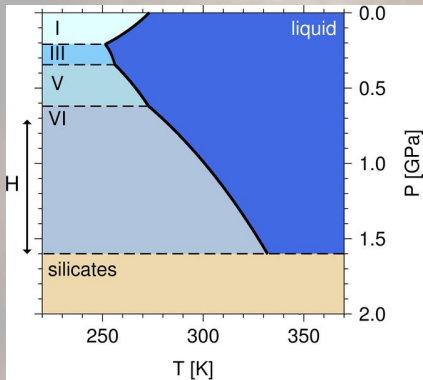
Titan

- ▶ partial differentiation:
Mol=0.3414
 - ▶ dense N₂-CH₄-Ar atmosphere
- possibly ongoing exchange?



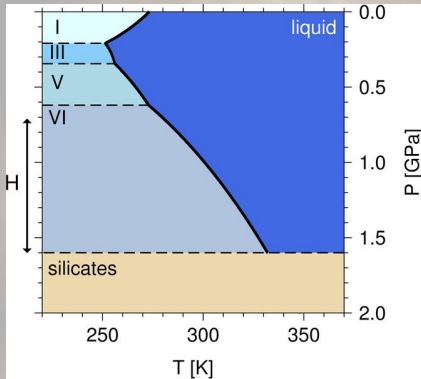
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- ▶ dehydrated mantle
+ thick H₂O layer
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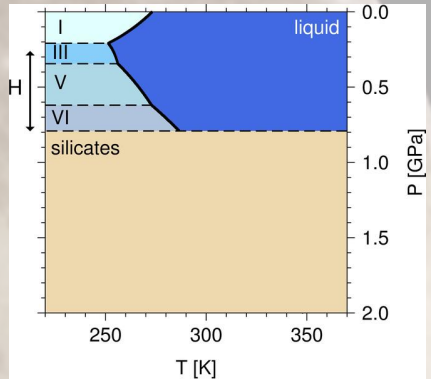
Ganymede

- ▶ dehydrated mantle + thick H₂O layer
- ▶ $H \gtrsim 400$ km: only ice VI



Titan

- ▶ core of hydrated silicates + thinner H₂O layer
- ▶ $H \sim 50\text{--}300$ km: ices VI, V, III?



Heat and material exchange

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- ▶ BUT: material exchange possible through convection
- ▶ HP ice dynamics governs heat extraction from the interior



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e.g. *Kirk & Stevenson (1987)*, *Grasset & Sotin (1996)*, *Showman & Malhotra (1997)*, *Tobie + (2005)*, *Fortes + (2007)*, *Grindrod + (2008)*, *Bland + (2009)*, *Noack + (2016)*

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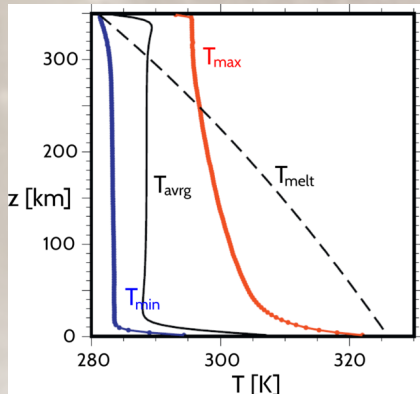
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- a) $T = T_m$, heat transfer through HP ice not treated
- b) use of scaling laws based on the hot TBL instability
 - ▶ effect of ice phases transitions on convection:
 - a) linear stability analysis (*Bercovici +, 1986*; *Sotin & Parmentier, 1989*)
 - b) 2d convection in ice VI and VII mantle (*Journaux & Noack, 2015*)

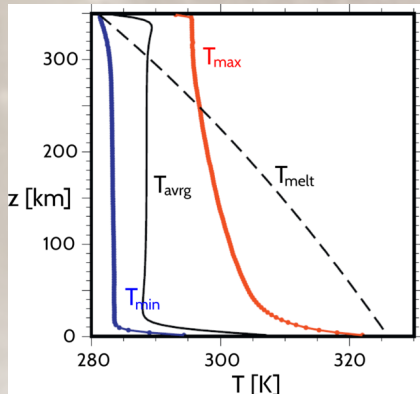
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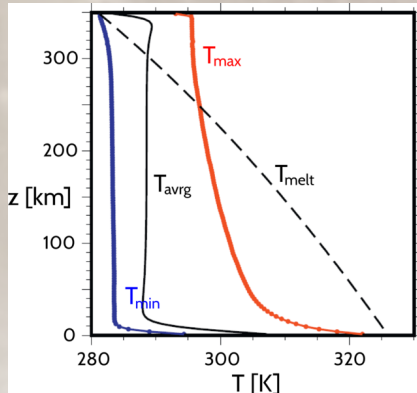
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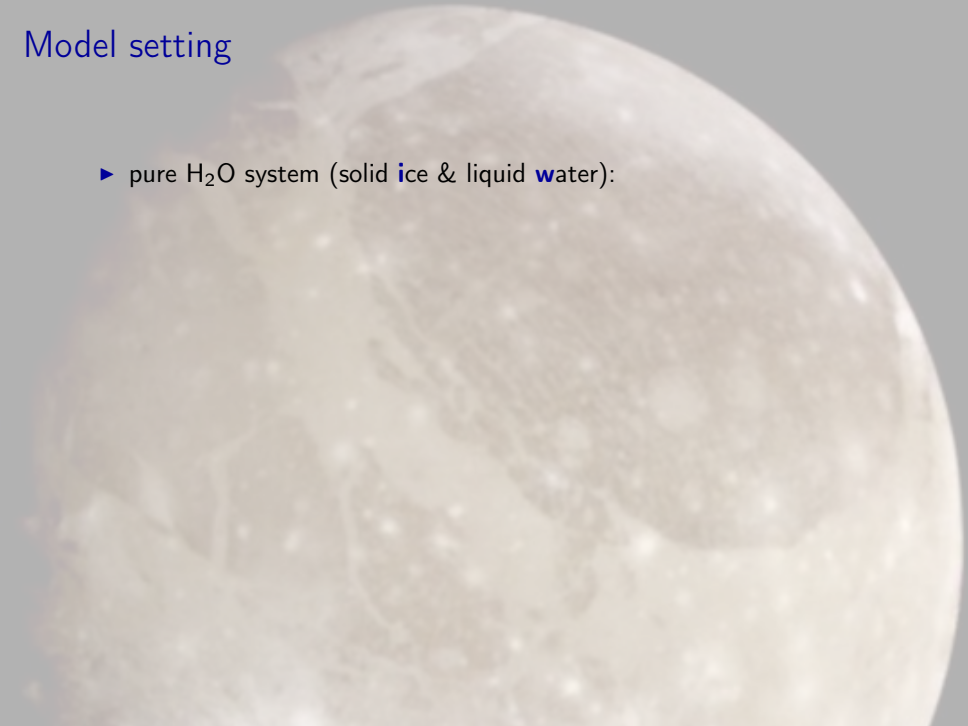
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- **two-phase mixture model** (*Bercovici +, 2001*; *Šrámek +, 2007*)



Model setting

- ▶ pure H₂O system (solid **i**ce & liquid **w**ater):



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a) $T_i < T^m$: one-phase material: **cold ice**

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▶ amount of ice: $1 - \phi$

▶ different material densities → **2 mass balances**

▶ phase separation → **2 linear momentum balances**

▶ thermal equilibrium between the two phases → **1 energy balance**

Governing equations

1. Thermal convection in ice - incompressible extended Boussinesq approximation

- ▶ mass balance of ice

$$\nabla \cdot \mathbf{v}_i = 0$$

- ▶ linear momentum balance of ice

$$0 = -\nabla \Pi - \rho_i \alpha_i T \mathbf{g} + \nabla \cdot \boldsymbol{\sigma}_i$$

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4. **Melting/freezing**

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2. Mixture of two incompressible components: $\xi_i \rightarrow \bar{\xi} := (1-\phi)\xi_i + \phi\xi_w$
3. $\mu_w \ll \mu_i$: $\bar{\sigma} \rightarrow (1-\phi)\sigma_i$, $T_w = T^m$: $\bar{\rho\alpha} \rightarrow (1-\phi)\rho_i\alpha_i$, $\bar{\rho c^P} \rightarrow (1-\phi)\rho_i c_i^P \rightarrow \rho_i c_i^P$
4. Melting/freezing \rightarrow **compressible mixture**

- ▶ mass balance of mixture

$$\nabla \cdot \bar{\mathbf{v}} = \frac{\Delta \rho}{\rho_i \rho_w} \Gamma_m$$

- ▶ linear momentum balance of ice

$$0 = -\nabla \Pi - (1-\phi)\rho_i \alpha_i T \mathbf{g} + \nabla \cdot (1-\phi)\boldsymbol{\sigma}_i$$

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$$\rho_i c_i^P \frac{\partial T}{\partial t} + \overline{\rho c^P \mathbf{v}} \cdot \nabla T - \rho_i T \overline{\alpha \mathbf{v}} \cdot \mathbf{g} = \nabla \cdot (\bar{k} \nabla T) + (1-\phi)\boldsymbol{\sigma}_i : \nabla \mathbf{v}_i$$

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- ▶ mass balance of water \rightarrow advection of porosity

$$\frac{\partial \phi}{\partial t} + \nabla \cdot (\phi \mathbf{v}_w) = \frac{\Gamma_m}{\rho_w}$$

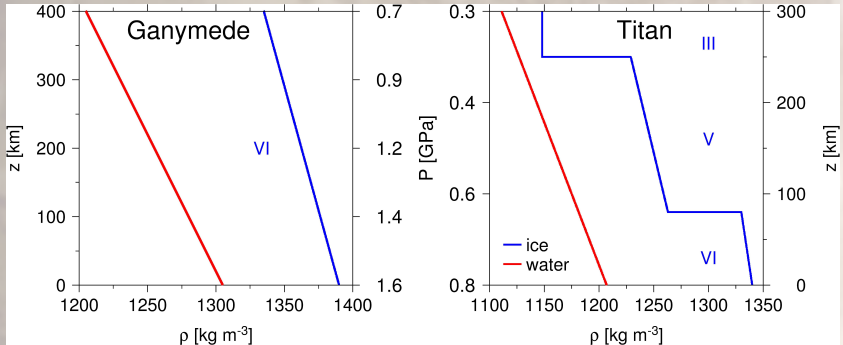
Material parameters

- ▶ ice permeable to water transport above a threshold porosity ϕ_c

$$K_\phi = \begin{cases} k_0(\phi - \phi_c)^2 & \phi \geq \phi_c \\ 0 & \phi < \phi_c \end{cases}$$

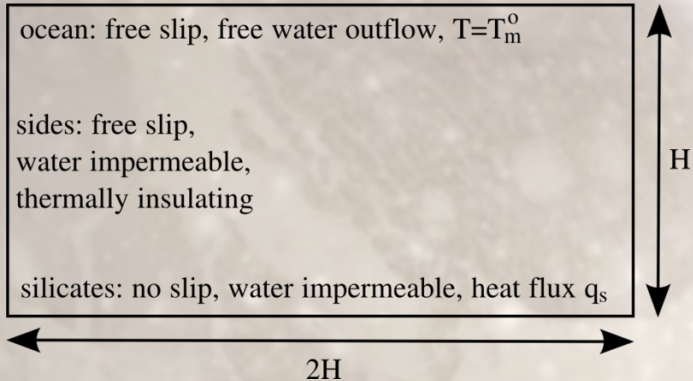
- ▶ temperature and depth dependent ice viscosity

$$\mu_i = \mu_0 \exp \left[\frac{Q}{R} \left(\frac{1}{T} - \frac{1}{T^m(z)} \right) \right]$$



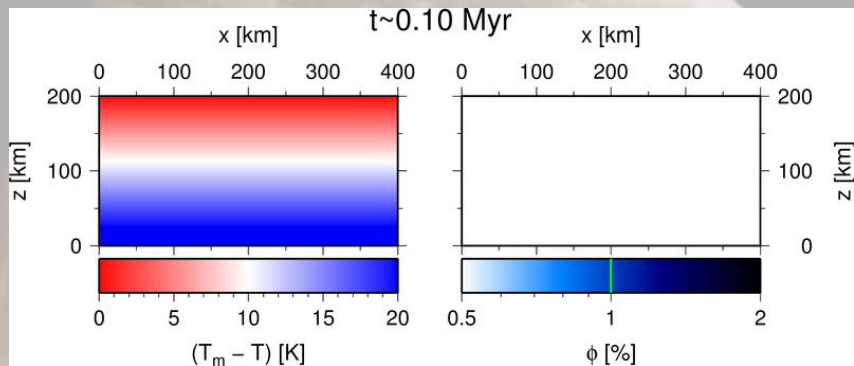
Implementation

- ▶ FEniCS (fenicsproject.org, Logg +, 2012; Alnaes +, 2015), MPI



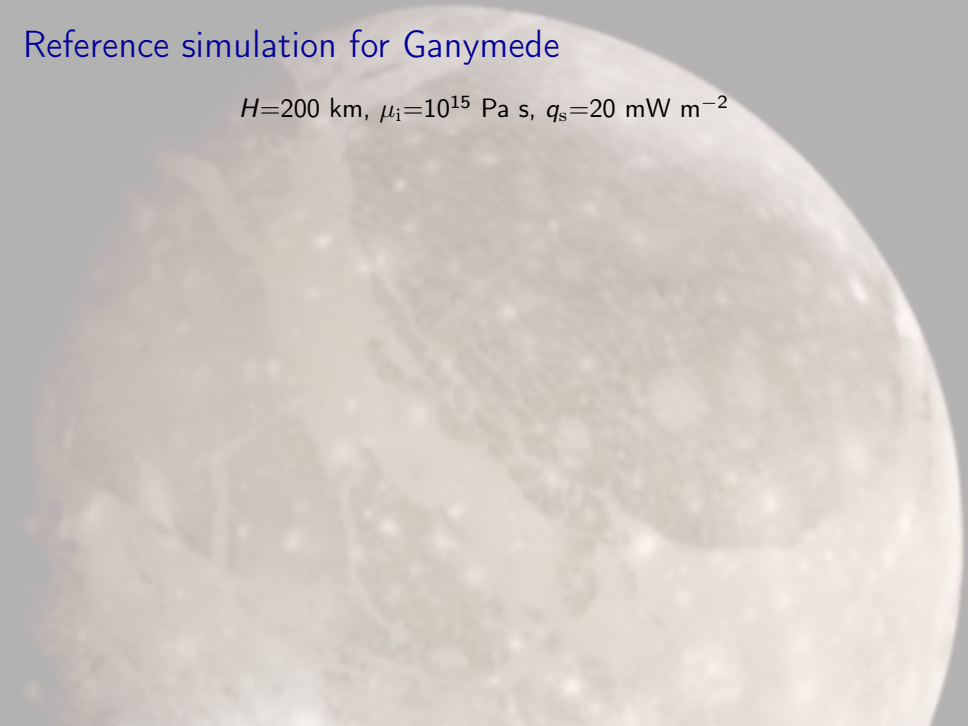
Reference simulation for Ganymede

$$H=200 \text{ km}, \mu_i=10^{15} \text{ Pa s}, q_s=20 \text{ mW m}^{-2}$$



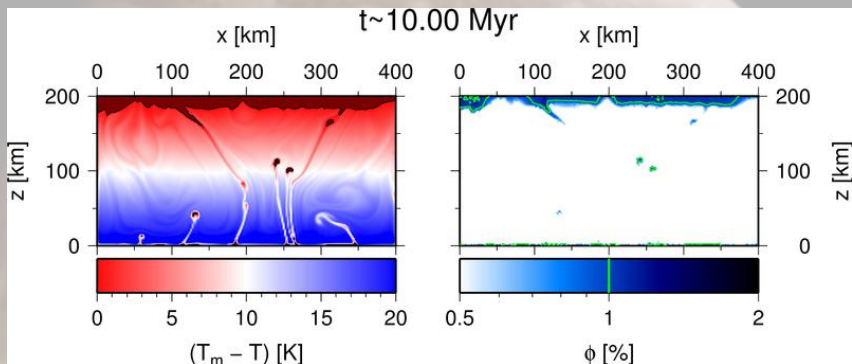
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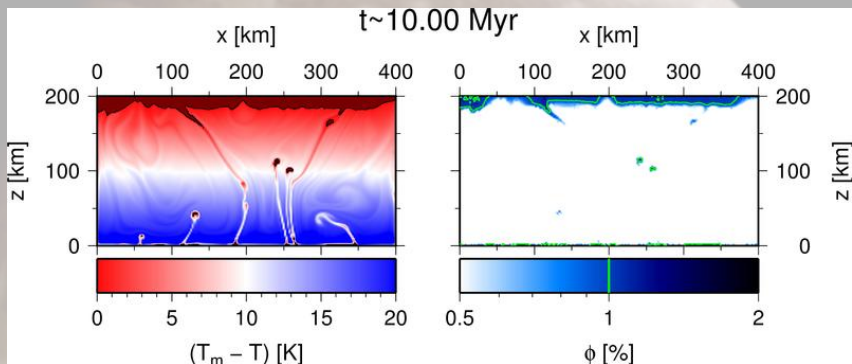
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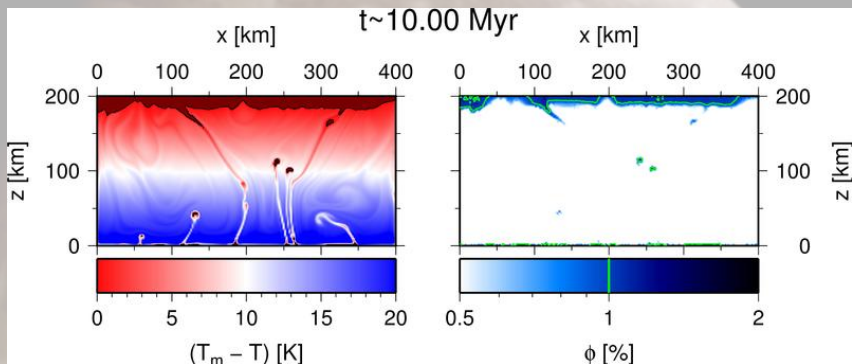
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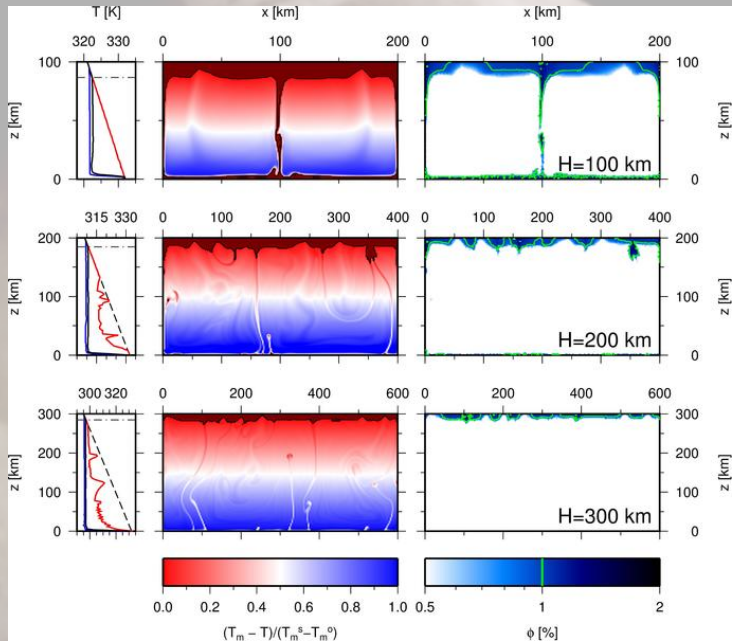
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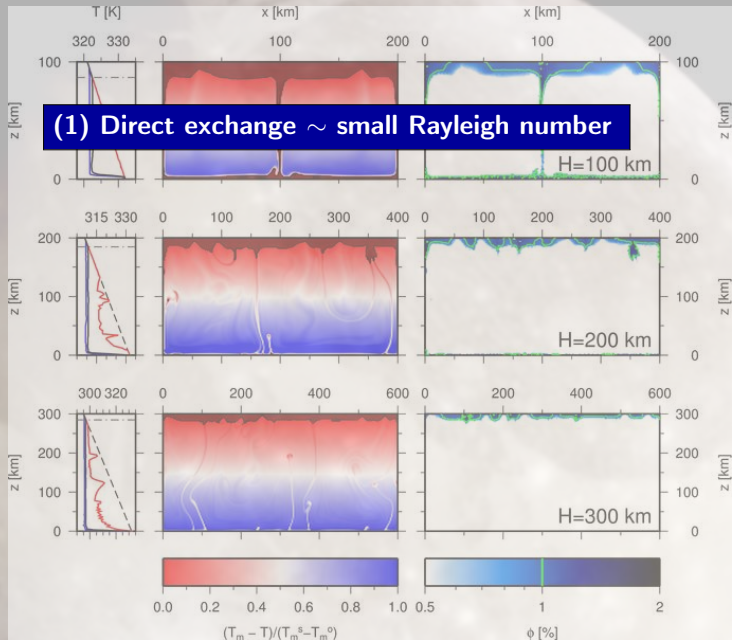


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- ▶ ocean: $T = T^m$, $\phi \sim \phi_c$: volatiles dissolution into the ocean

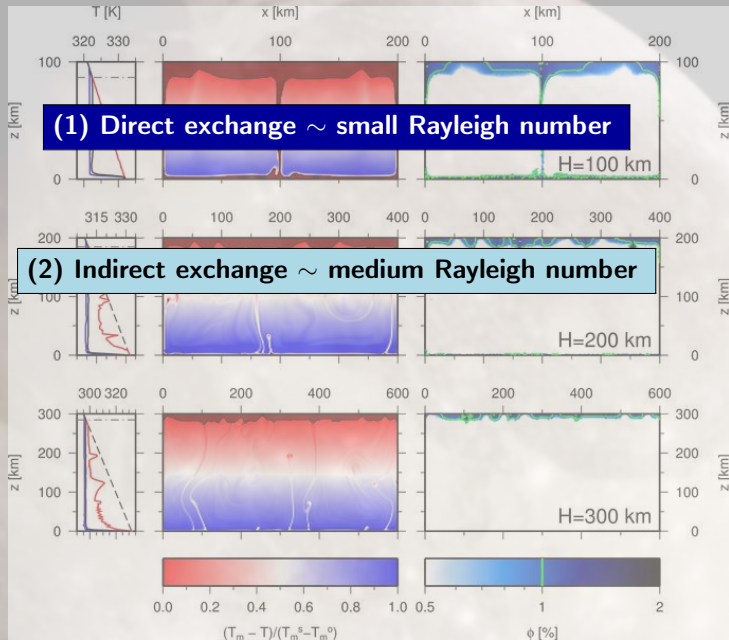
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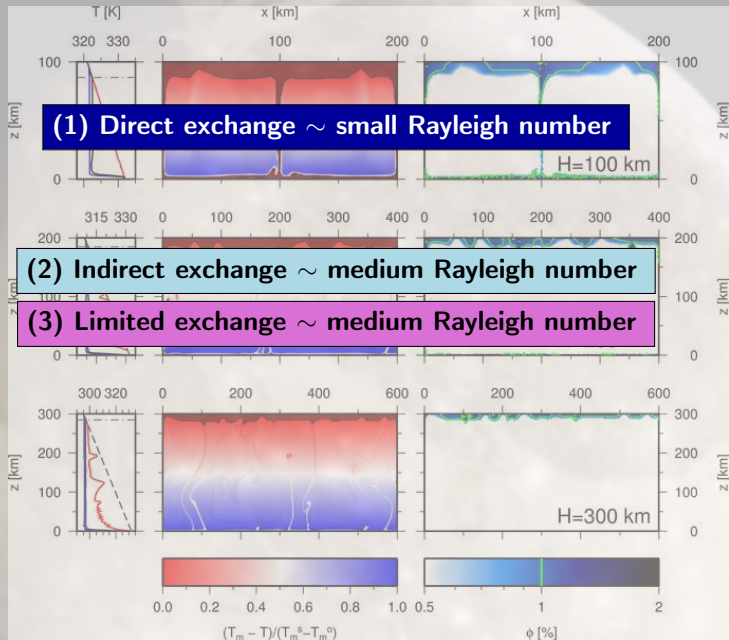
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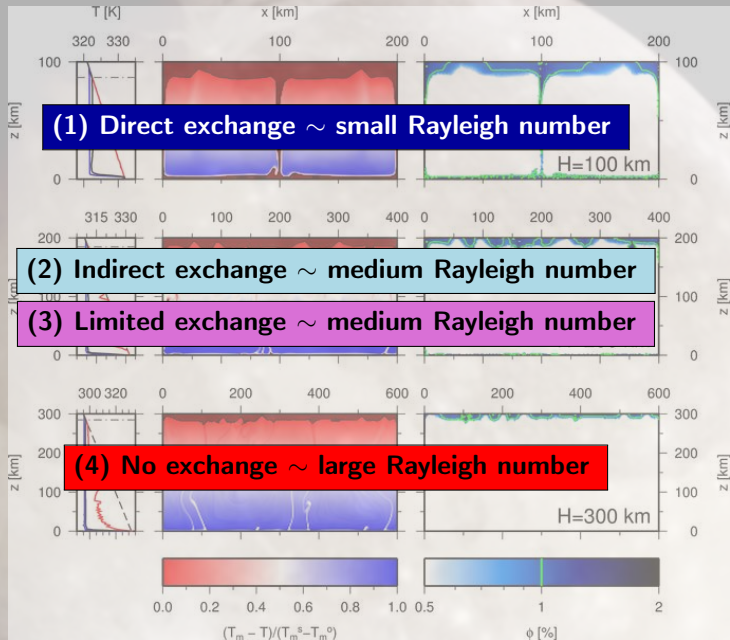
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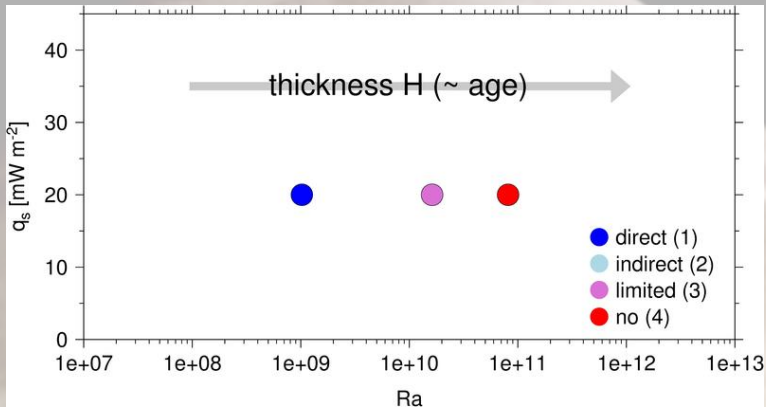
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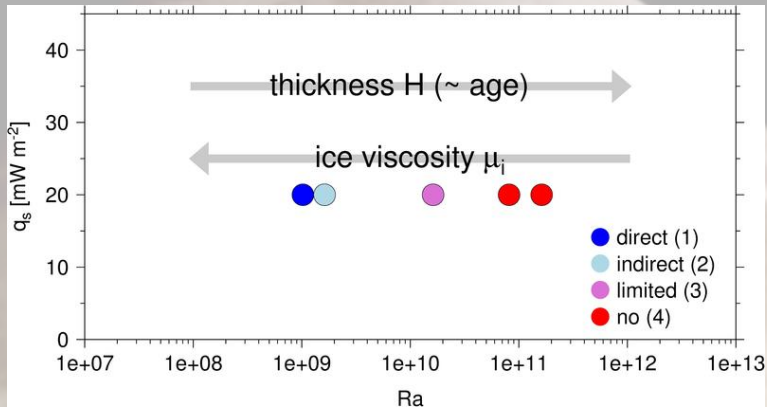


Results of parametric study



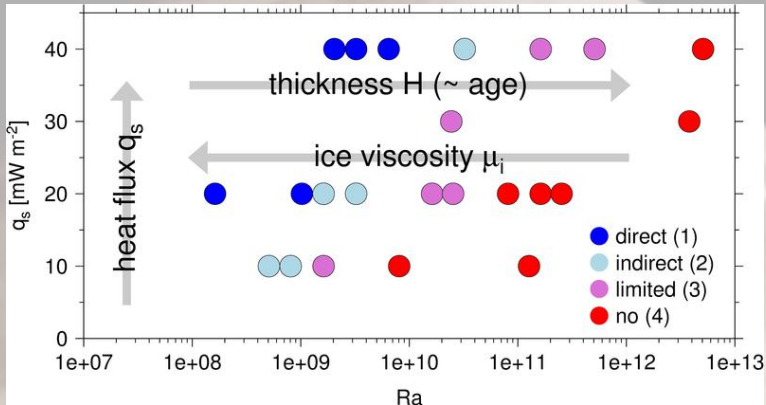
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Long-term evolution modeling



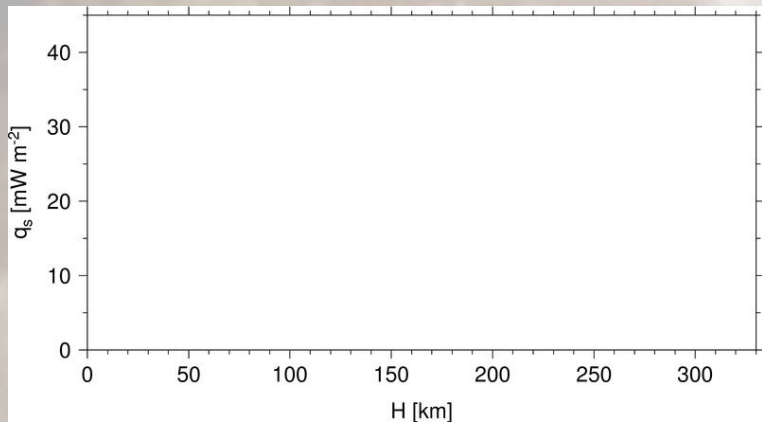
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- ▶ 1d thermo(-chemical) evolution model:
 - scaling laws for rocky interior, **HP ice layer**, ocean, ice I crust
 - q_s and H are part of solution
 - μ_i remains an input parameter

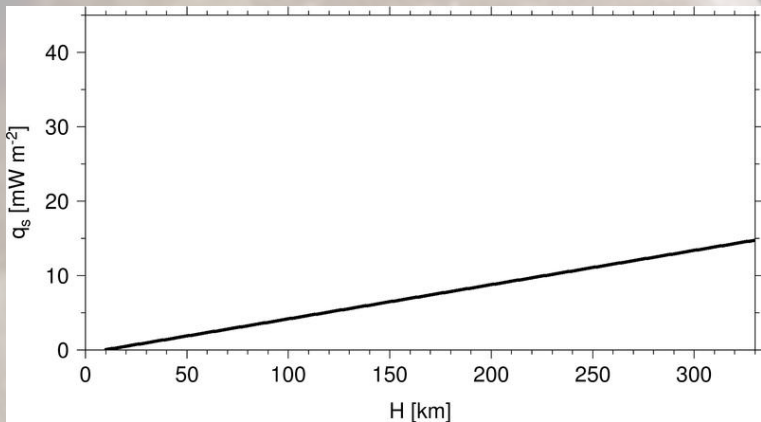
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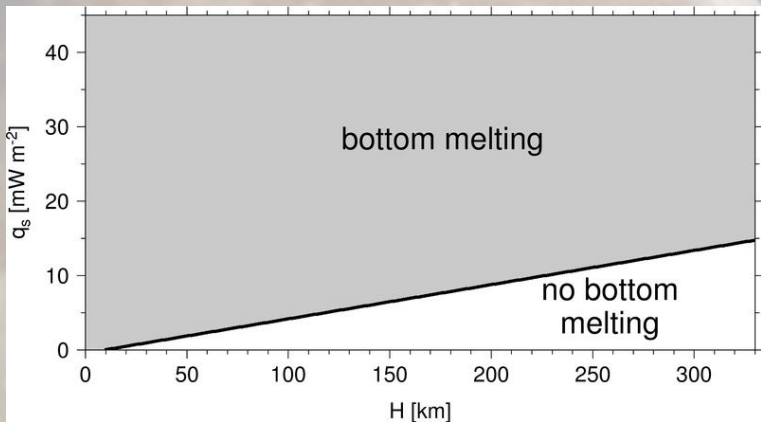
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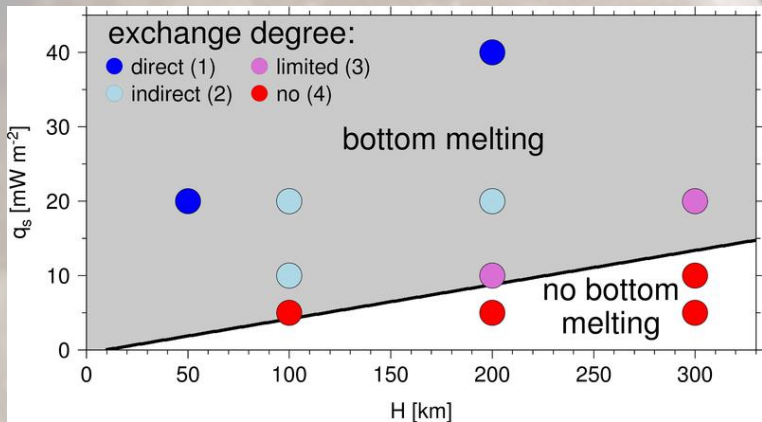
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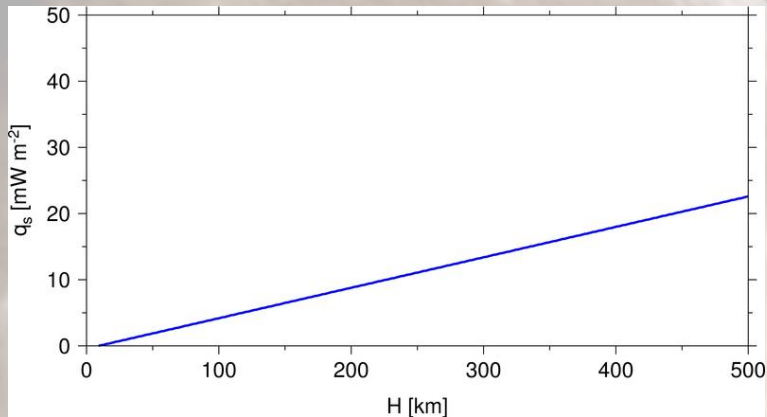
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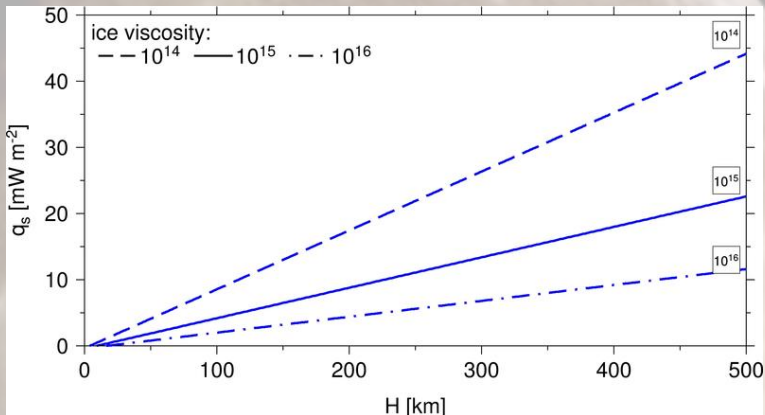
Present-day bottom melting: Ganymede vs Titan

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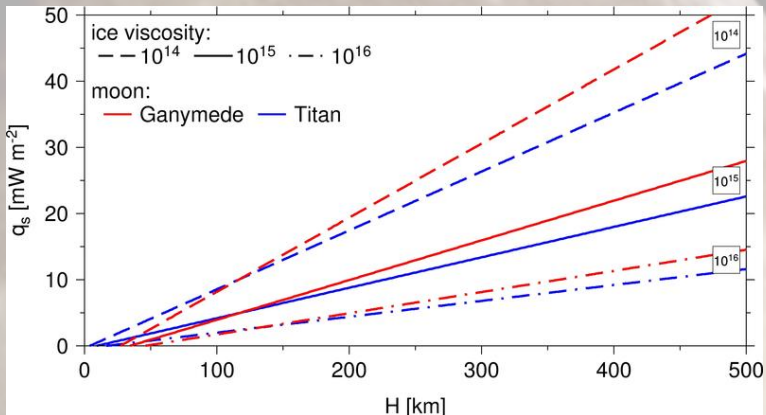
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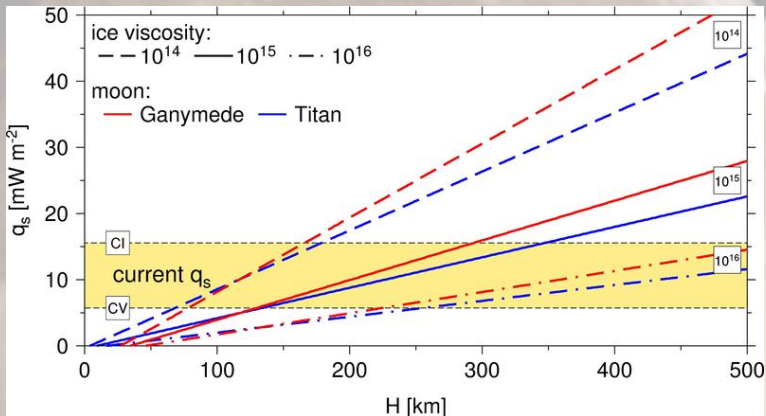
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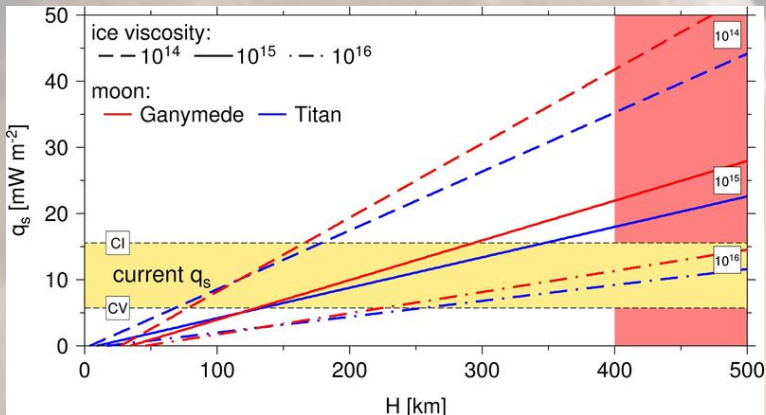
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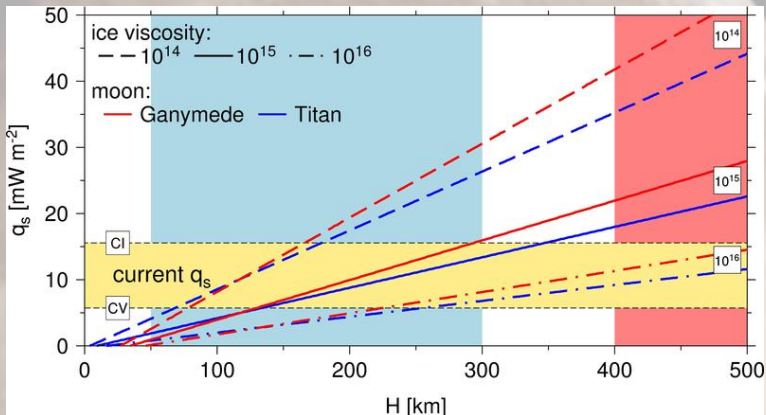
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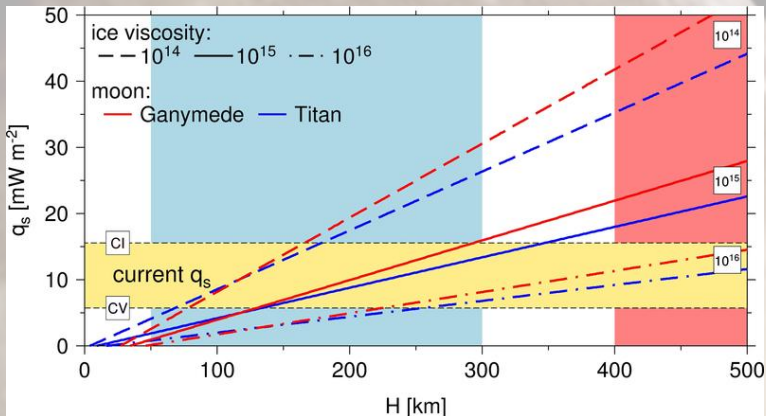
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→ Titan: melting at silicates interface and leaching of ^{40}Ar ?

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juice

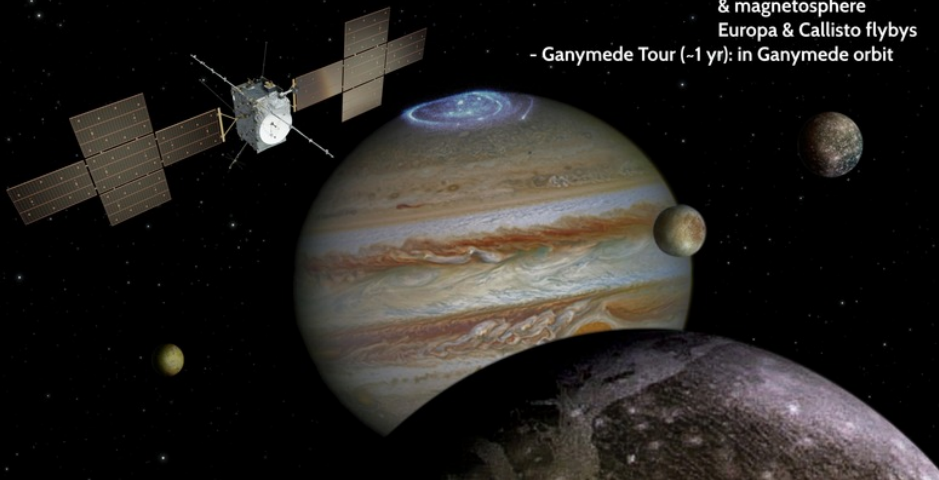


→ JUPITER ICY MOONS EXPLORER

Exploring the emergence of habitable worlds around gas giants

Two mission phases:

- Jupiter Tour (~2.5 yr): Jovian atmosphere & magnetosphere
Europa & Callisto flybys
- Ganymede Tour (~1 yr): in Ganymede orbit



Dragonfly



- rotocraft lander
- dozens of locations
- launch 2026, arrival 2034



Search for prebiotic chemical processes
common on both Titan and Earth ...

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Thank you for your
attention!

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